

THE 2019 CCNU-USTC JUNIOR COSMOLOGY SYMPOSIUM

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SHAFT POTENTIAL INFLATION

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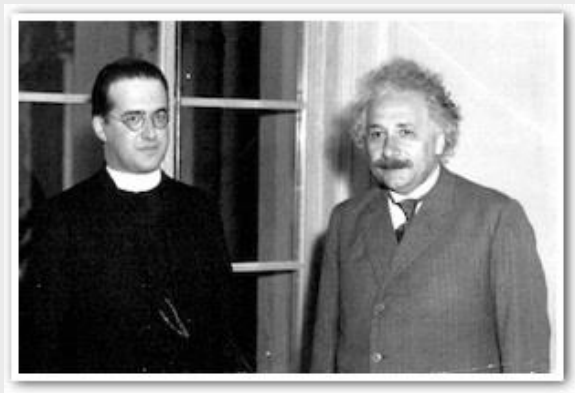
OUTLINE

- ❑ SHORT HISTORY OF UNIVERSE
- ❑ INFLATION'S BASICS
- ❑ RESEARCH DISCUSSION
- ❑ SUMMARY OF TALK

SHORT HISTORY OF UNIVERSE

BIG BANG

- **First general relativistic models by Einstein Predicted a contracting universe (1917)**
- **George Lemaitre was the first who proposed a "primeval atom" theory, later know as Big Bang Theory.**
- **Cosmic Microwave Background (CMB) discovered in 1964 by radio astronomers Arno Penzias and Robert Wilson.**
- **CMB caused most cosmologists to accept "Big Bang".**



SOLUTION WITH INFLATION

➤ FLATNESS:

Inflation pushes universe towards flatness.

Longer Inflation \Rightarrow Flatter Universe

➤ HORIZON:

Observable universe was extremely small before inflation \Rightarrow all regions could be in contact.

➤ RELIC:

Inflation massively dilutes any relics \Rightarrow we should not observe them today.



TACHYONIC FIELD INFLATION

In flat FRW metric, Friedmann equation is

$$H^2 = \frac{8\pi}{3M_p^2}(\rho_\phi + \rho_\gamma),$$

The energy density and pressure of tachyon scalar field are defined as follows

$$\rho_\phi = \frac{V(\phi)}{\sqrt{1 - \dot{\phi}^2}}, \quad p_\phi = -V(\phi)\sqrt{1 - \dot{\phi}^2}.$$

TACHYONIC FIELD INFLATION

The second conservation equation provided as

$$3H(1 + R)\dot{\phi} = -\frac{V'(\phi)}{V(\phi)}, \quad \text{where } \ddot{\phi} \ll \left(3H + \frac{\Gamma}{V}\right)\dot{\phi},$$

The slow-roll parameters can also be deduced in the form of scalar field and thermalization according to the tachyon field,

$$\epsilon = \frac{M_p^2}{2(1 + R)} \left(\frac{V'^2}{V^3}\right), \quad \eta = \frac{M_p^2}{(1 + R)} \left(\frac{V''}{V^2}\right).$$

SHAFT INFLATIONARY PARAMETERS

The generalized form of shaft potential is

$$V(\phi) = \frac{M_p^4 \phi^{2n-2}}{(\phi^n + m^n)^{2-\frac{2}{n}}},$$

The number of e-folds can be calculated with the help of tachyon field and under shaft potential

$$N = \frac{M_p^2}{(2-2n)m^n} \int_{\phi_i}^{\phi_f} \phi^{-1+2n} (m^n + \phi^n)^{-1+\frac{2}{n}} (1+R) d\phi.$$

SHAFT INFLATIONARY PARAMETERS

Tensor to scalar ratio with the help of shaft potentiation with respect to techyonic field

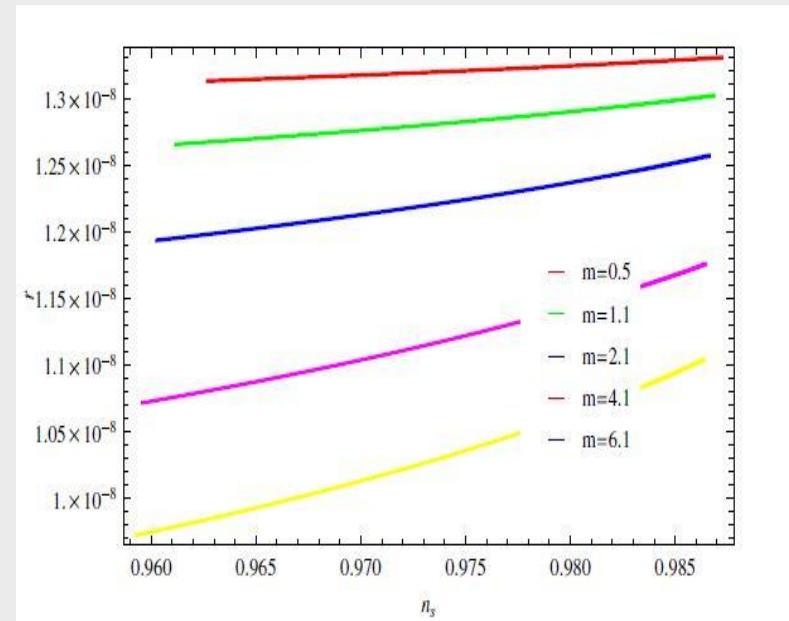
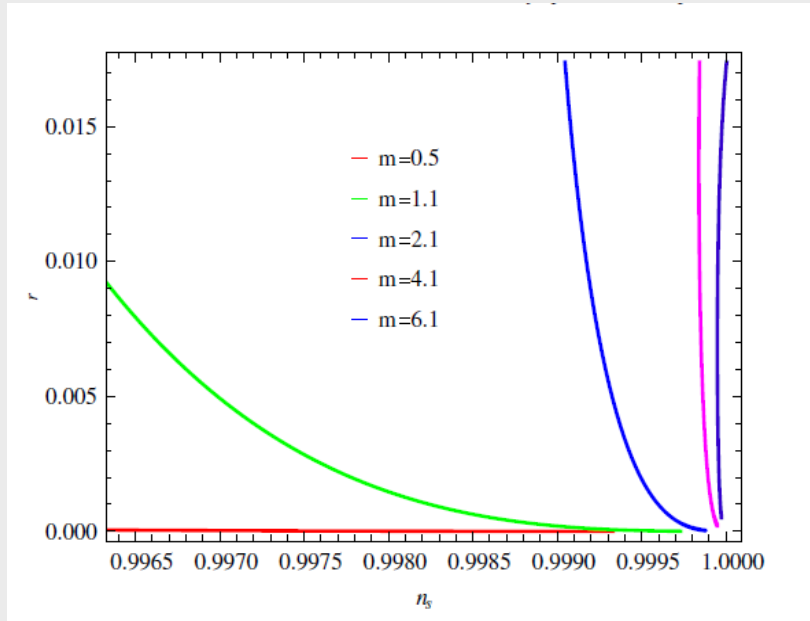
$$r = \frac{5123^{-\frac{-3+q}{2(-4+q)}} am^{3n} (-1+n)^3 \phi^{4-6n-q} (m^n + \phi^n)^{3-\frac{6}{n}}}{C_\gamma M_p^8 (1+R)^{11/2}} \times \left(\frac{am^{2n} M_p^3 (-1+n)^2 \phi^{-1-q}}{C_\gamma (1+R)^2 (m^n + \phi^n)^2 (M_p^4 \phi^{-2+2n} (m^n + \phi^n)^{-2+\frac{2}{n}})^{3/2}} \right)^{-1+\frac{1}{-4+q}}$$

SHAFT INFLATIONARY PARAMETERS

However, the spectral index and its running attained the values

$$n_s = 1 + \frac{1}{M_p^2(-4+q)(1+R)} (2m^n(-1+n)\phi^{-2n}(m^n + \phi^n)^{-2/n} \\ \times (m^n(10 - 17n - 2q + 5nq) + (-7 + 2n(-3 + q) + 3q)\phi^n)),$$
$$\alpha_s = \frac{1}{M_p^2(-4+q)(1+R)} (4m^{2n}(-1+n)^2\phi^{-4n}(m^n + \phi^n)^{-4/n} \\ \times (2m^{2n}n(10 - 17n - 2q + 5nq) + m^n(-4(-5 + q) + n(-21 + 9q \\ + 4n(-10 + 3q)))\phi^n + (2 + n)(-7 + 2n(-3 + q) + 3q)\phi^{2n})).$$

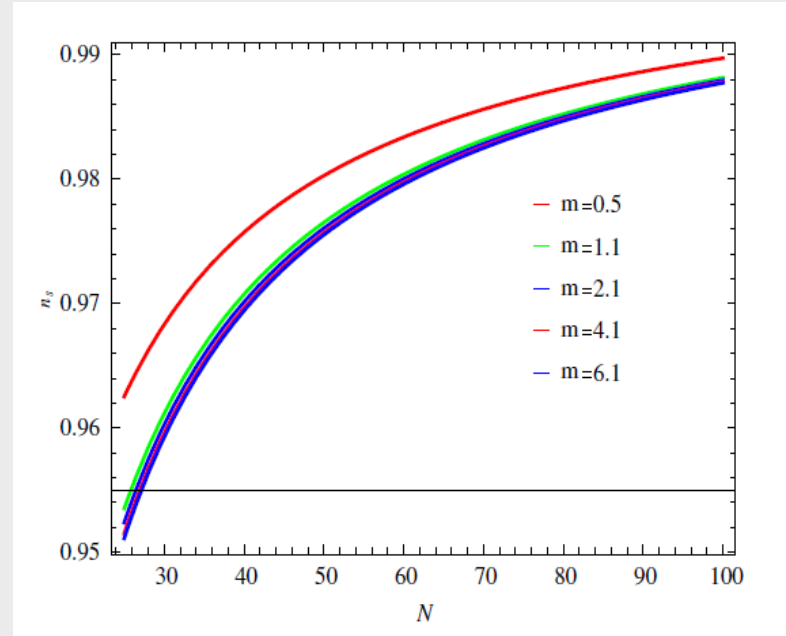
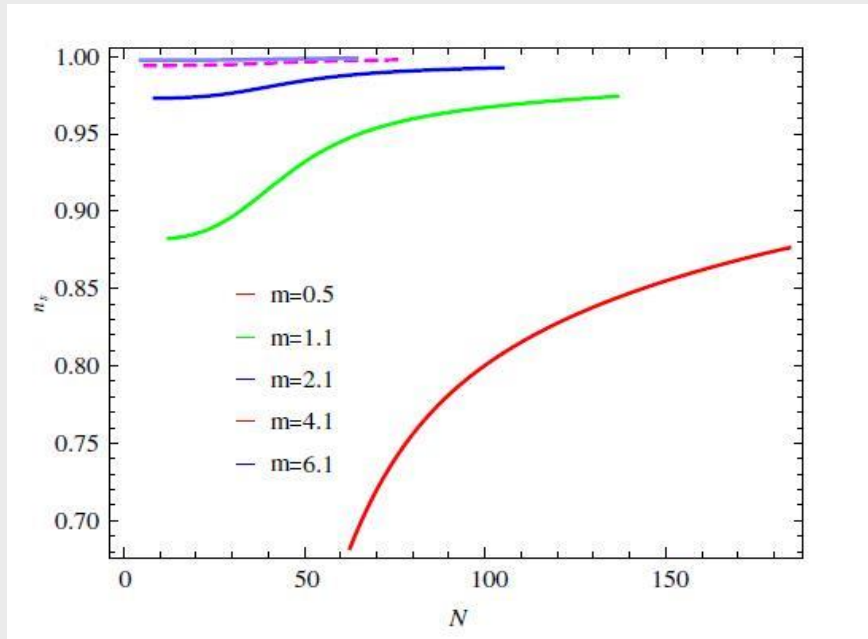
GRAPHICAL REPRESENTATION



TENSOR-SCALAR RATIO

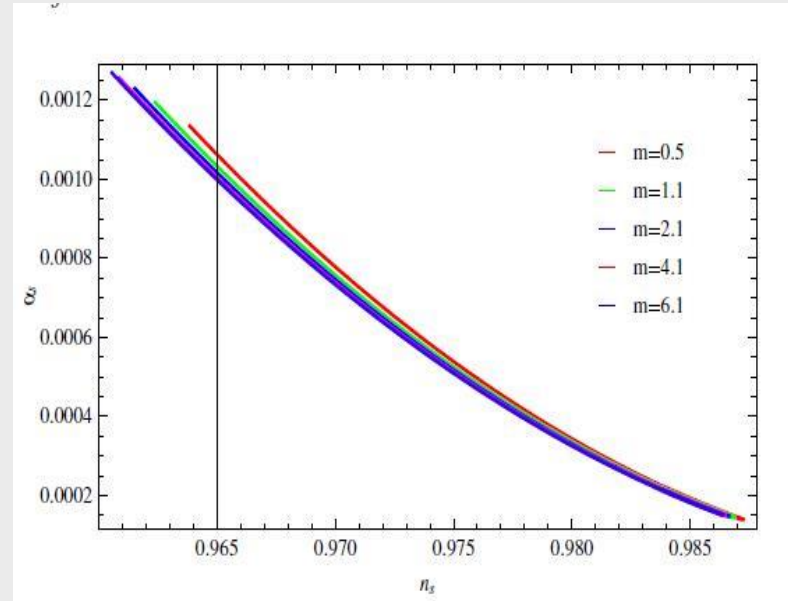
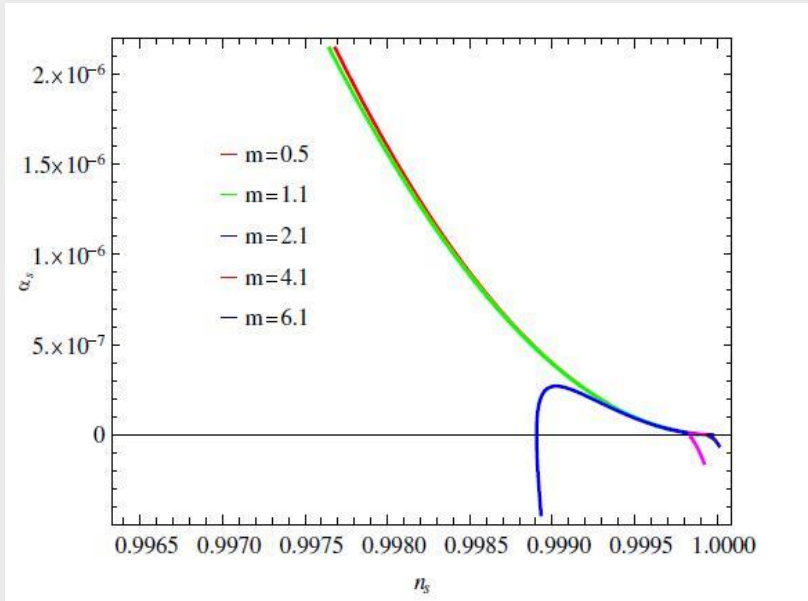


GRAPHICAL REPRESENTATION

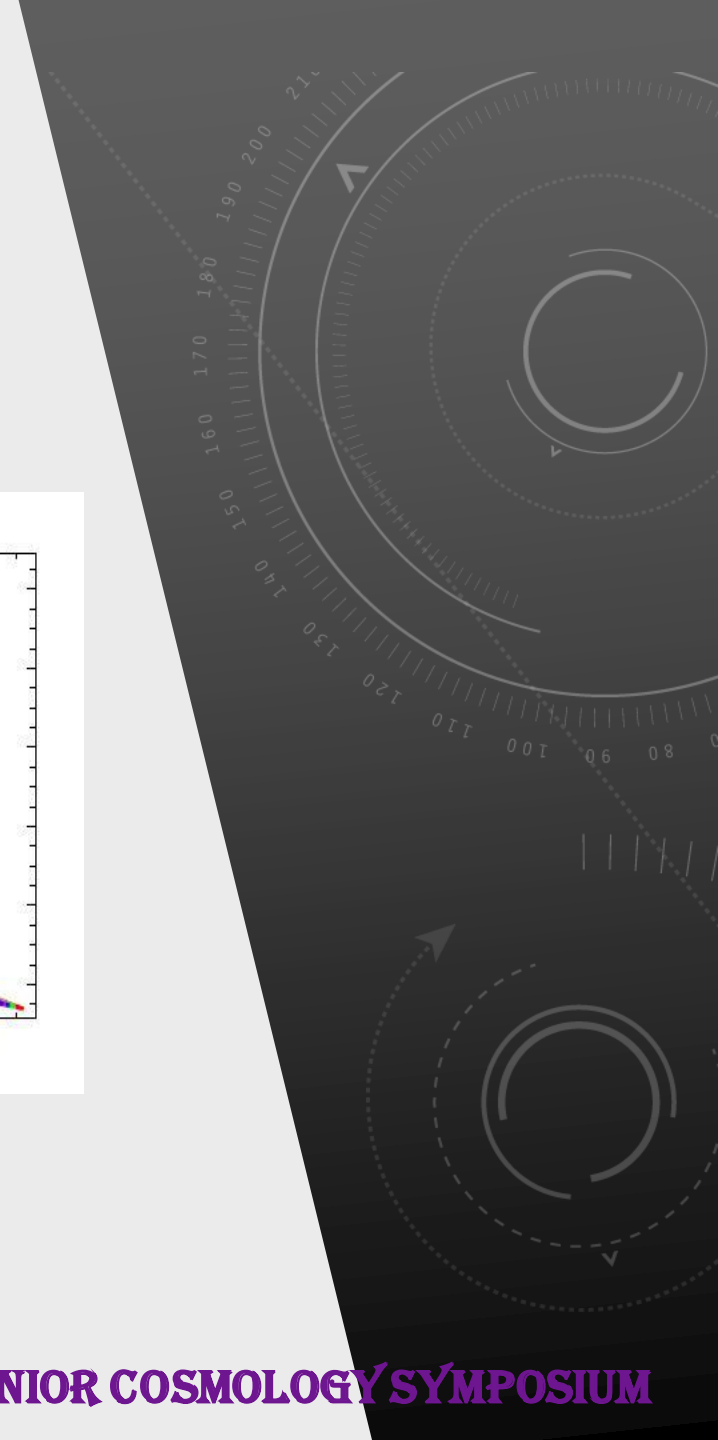


SPECTRAL INDEX

GRAPHICAL REPRESENTATION



RUNNING OF SPECTRAL INDEX



SUMMARY

- ❖ **Examined the effect of shaft potential on the warm tachyon fields inflationary scenario.**
- ❖ **We have considered the scalar field decay rate which is proportional to the temperature**
- ❖ **Investigated the inflationary parameters in both strong and weak dissipative regimes.**
- ❖ **Showing consistency with Observational data.**

THANK YOU